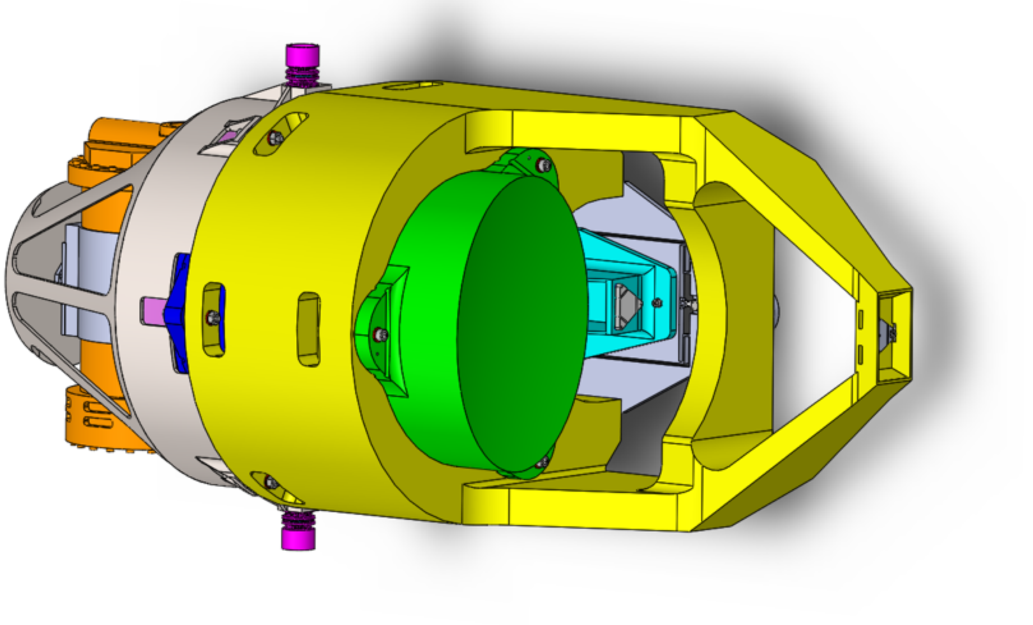
Optical Telescopes for the L3/LISA Space-Based Gravitational Wave Observatory



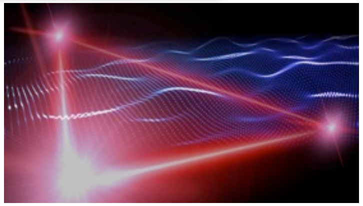
Jeff Livas for the US LISA Telescope Team

NASA Goddard Space Flight Center

Greenbelt, MD 20771

Nov 2017

Telescope Team



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- Justin WARD, Joseph IVANOV, Alex MILLER

EDGE Space Systems (thermal): Angel DAVIS

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University of Florida:

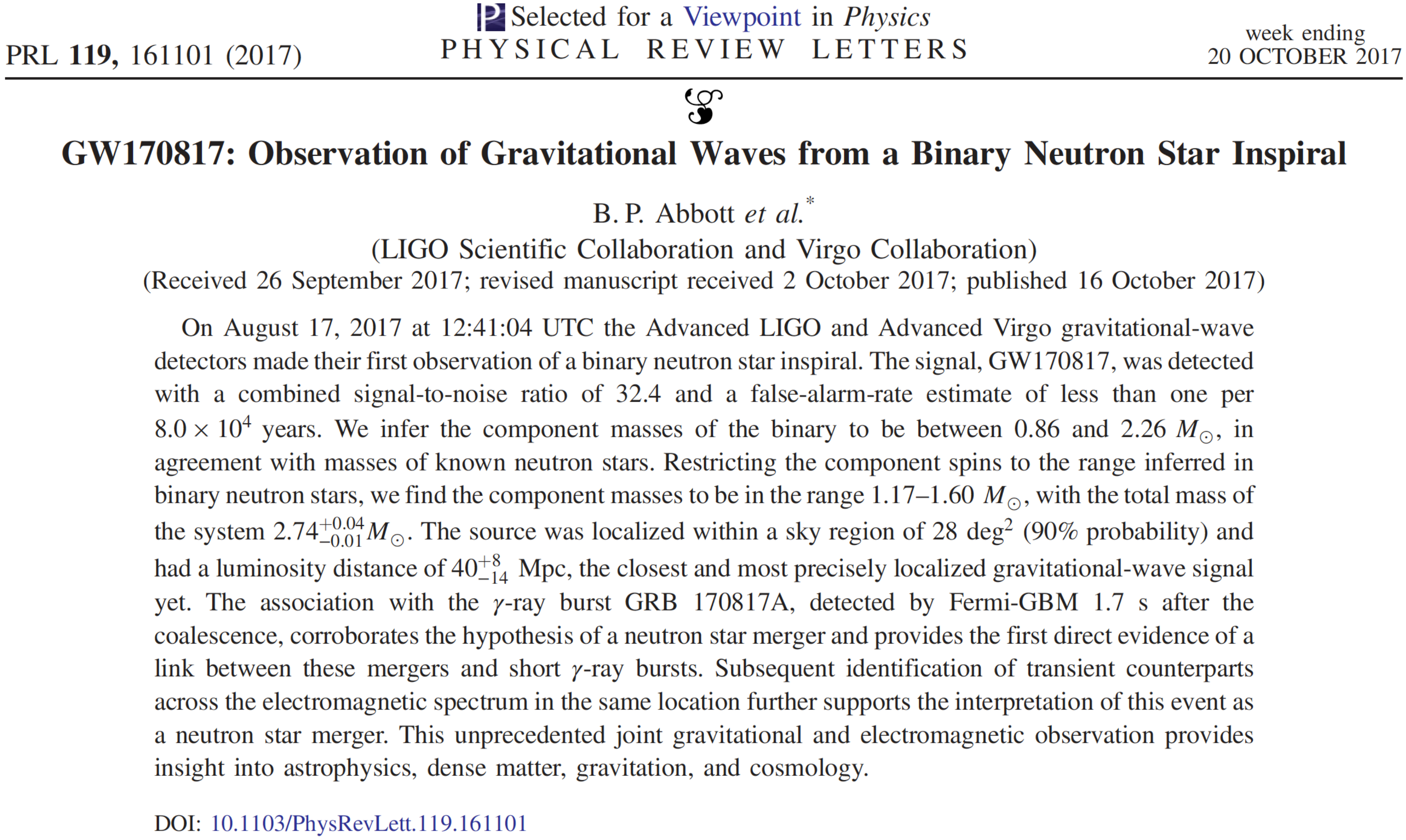
- Professor Guido MUELLER's group

This work was supported by NASA grants 11-SAT11-0027 and 14-SAT14-0014.

Outline

- Mission Context and Science
- Measurement Principles
- Telescope Description
- Challenges
- Summary

MISSION CONTEXT AND SCIENCE



Selected for a Viewpoint in *Physics*

PRL **119,** 161101 (2017) PHYSICAL REVIEW LETTERS week ending 20 OCTOBER 2017

GW170817: Observation of Gravitational Waves from a Binary Neutron Star Inspiral

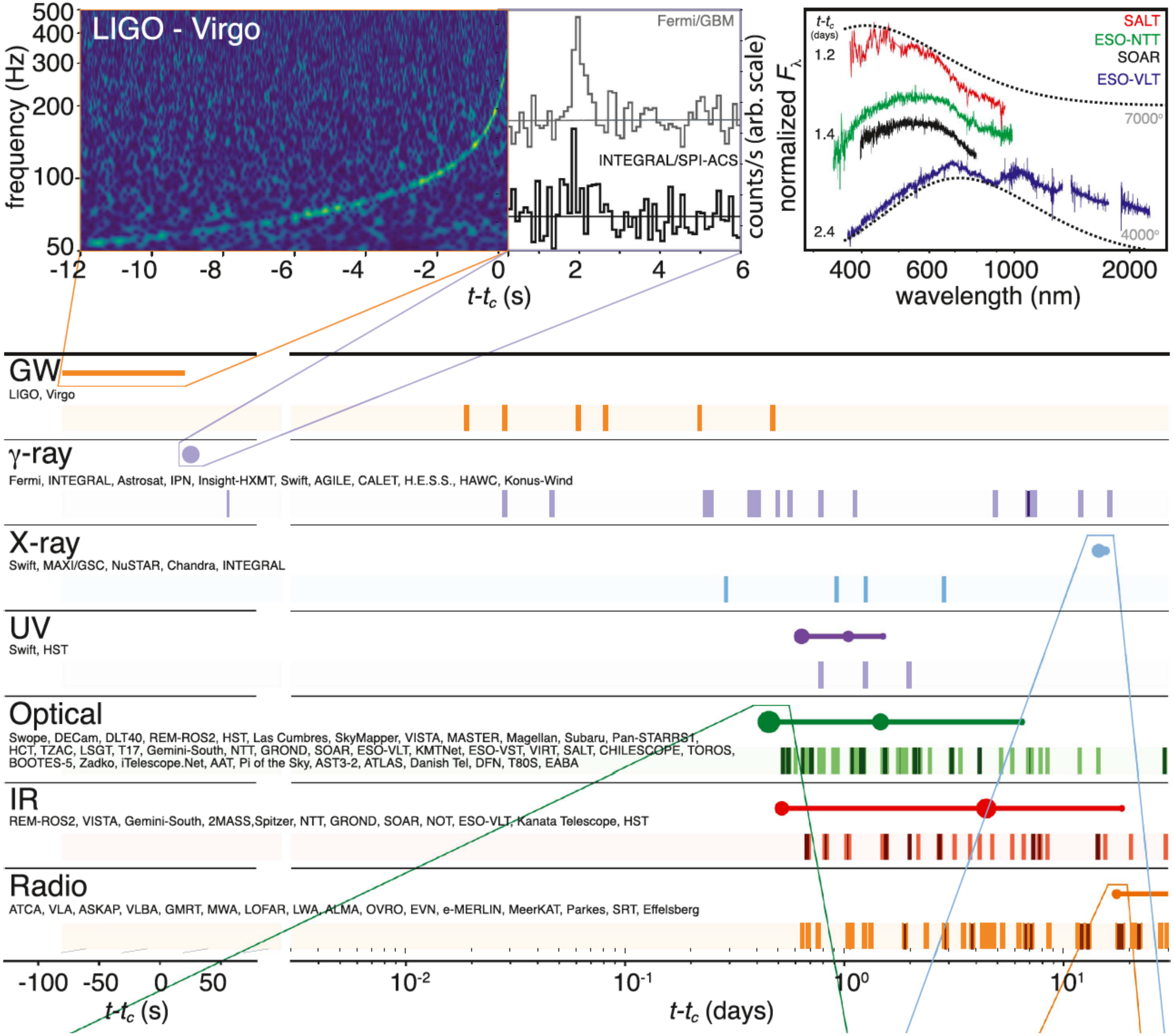
B. P. Abbott *et al.*[*]

(LIGO Scientific Collaboration and Virgo Collaboration)

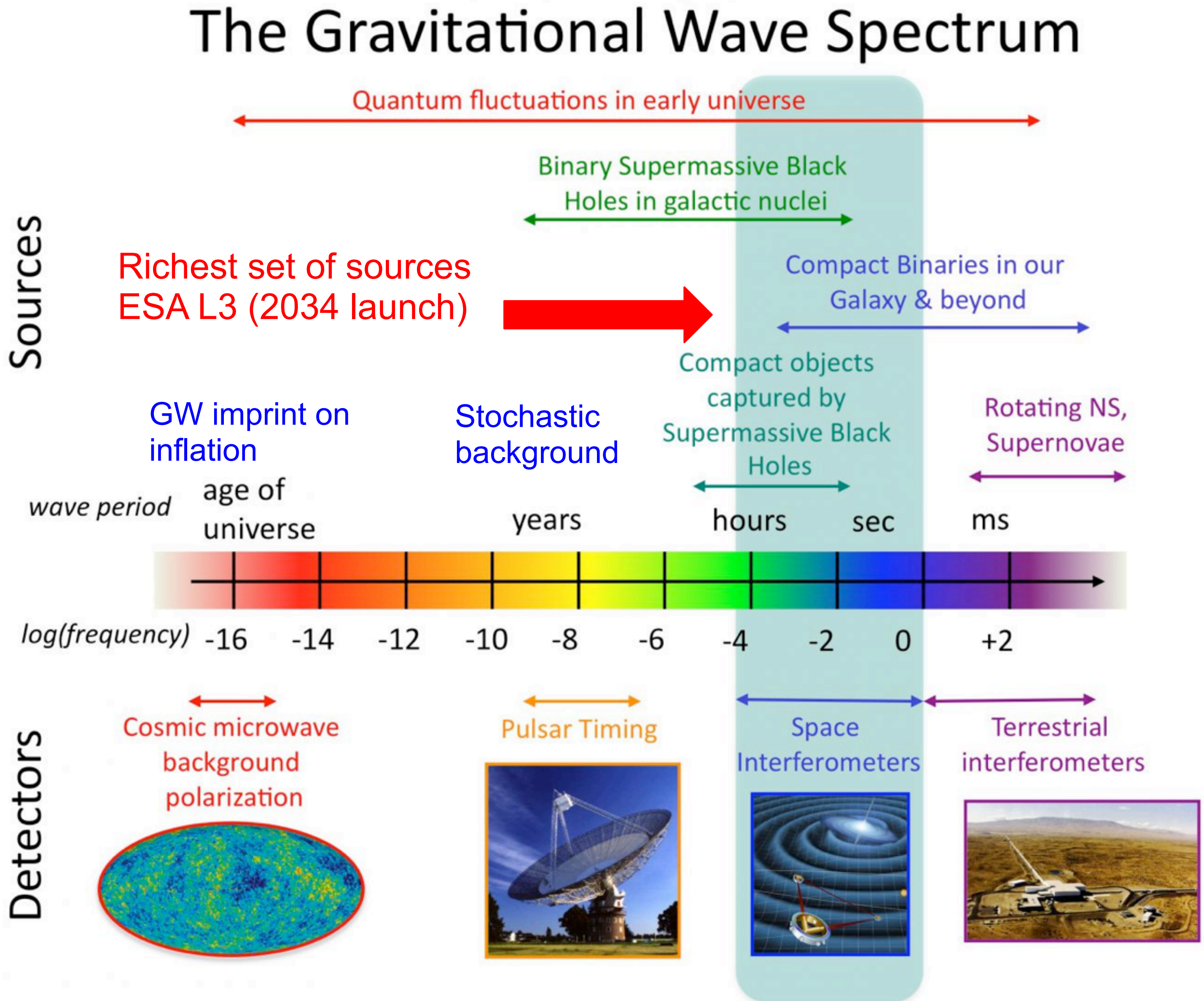
(Received 26 September 2017; revised manuscript received 2 October 2017; published 16 October 2017)

On August 17, 2017 at 12:41:04 UTC the Advanced LIGO and Advanced Virgo gravitational-wave detectors made their first observation of a binary neutron star inspiral. The signal, GW170817, was detected with a combined signal-to-noise ratio of 32.4 and a false-alarm-rate estimate of less than one per 8.0×10^4 years. We infer the component masses of the binary to be between 0.86 and 2.26 $M_\odot$, in agreement with masses of known neutron stars. Restricting the component spins to the range inferred in binary neutron stars, we find the component masses to be in the range 1.17–1.60 $M_\odot$, with the total mass of the system $2.74^{+0.04}_{-0.01} M_\odot$. The source was localized within a sky region of 28 deg^2 (90% probability) and had a luminosity distance of 40^{+8}_{-14} Mpc, the closest and most precisely localized gravitational-wave signal yet. The association with the γ-ray burst GRB 170817A, detected by Fermi-GBM 1.7 s after the coalescence, corroborates the hypothesis of a neutron star merger and provides the first direct evidence of a link between these mergers and short γ-ray bursts. Subsequent identification of transient counterparts across the electromagnetic spectrum in the same location further supports the interpretation of this event as a neutron star merger. This unprecedented joint gravitational and electromagnetic observation provides insight into astrophysics, dense matter, gravitation, and cosmology.

DOI: 10.1103/PhysRevLett.119.161101


LIGO - Virgo
frequency (Hz)
500
400
300
200
100
50
-12
-10
-8
-6
-4
-2
0
2
4
6
t-t_c (s)
Fermi/GBM
INTEGRAL/SPI-ACS
counts/s (arb. scale)
normalized F_λ
t-t_c (days)
1.2
1.4
2.4
SALT
ESO-NTT
SOAR
ESO-VLT
7000°
4000°
400
600
1000
2000
wavelength (nm)
GW
LIGO, Virgo
γ-ray
Fermi, INTEGRAL, Astrosat, IPN, Insight-HXMT, Swift, AGILE, CALET, H.E.S.S., HAWC, Konus-Wind
X-ray
Swift, MAXI/GSC, NuSTAR, Chandra, INTEGRAL
UV
Swift, HST
Optical
Swope, DECam, DLT40, REM-ROS2, HST, Las Cumbres, SkyMapper, VISTA, MASTER, Magellan, Subaru, Pan-STARRS1, HCT, TZAC, LSGT, T17, Gemini-South, NTT, GROND, SOAR, ESO-VLT, KMTNet, ESO-VST, VIRT, SALT, CHILESCOPE, TOROS, BOOTES-5, Zadko, iTelescope.Net, AAT, Pi of the Sky, AST3-2, ATLAS, Danish Tel, DFN, T80S, EABA
IR
REM-ROS2, VISTA, Gemini-South, 2MASS,Spitzer, NTT, GROND, SOAR, NOT, ESO-VLT, Kanata Telescope, HST
Radio
ATCA, VLA, ASKAP, VLBA, GMRT, MWA, LOFAR, LWA, ALMA, OVRO, EVN, e-MERLIN, MeerKAT, Parkes, SRT, Effelsberg
-100
-50
0
50
t-t_c (s)
10^{-2}
10^{-1}
10^{0}
10^{1}
t-t_c (days)

Why is this important?



Richest set of sources
ESA L3 (2034 launch)

GW imprint on inflation

Stochastic background

BICEP-2/WMAP/Planck

Detection 2018-20?

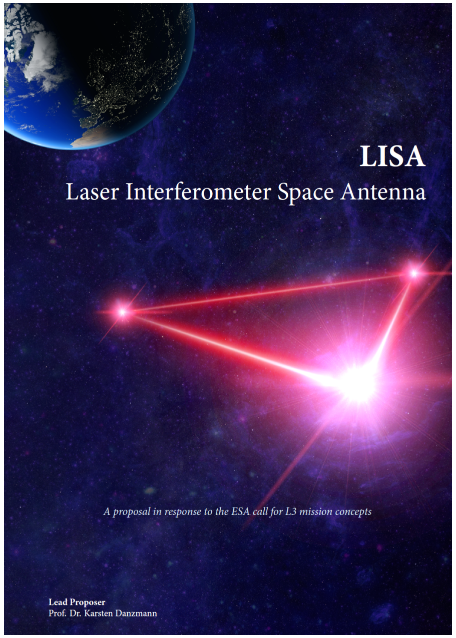
Detection 2015!

Image credit: NASA

ESA/NASA Activities

- Phase A to start early 2018:
  - Follows selection by SPC earlier this year
  - Intended to be competitive industrial study
  - 18 month duration
  - ESA Study Office has been established
  - Science Study Team has been established
  - US team also assembled to address decadal survey

  https://www.elisascience.org/files/publications/LISA_L3_20170120.pdf



LISA

Laser Interferometer Space Antenna

A proposal in response to the ESA call for L3 mission concepts

Lead Proposer
Prof. Dr. Karsten Danzmann

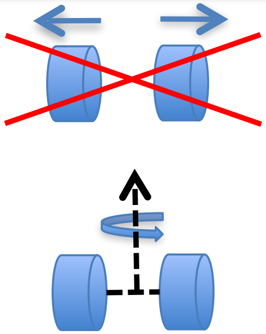
- GSFC plans:
  - Plan to produce a Breadboard by 2022
  - Currently iterating through optical/structural/thermal design
  - Other technologies also under development

https://lisa.nasa.gov/

MEASUREMENT PRINCIPLES

Measurement Challenge

- Lowest order radiator is a quadrupole
  - Dipole radiation forbidden by conservation of momentum
  - Simplest quadrupole: a “dumbell”

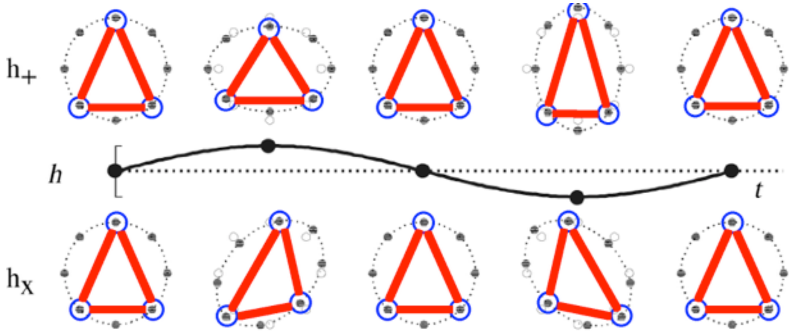


h_+ Polarization

h_x Polarization

- What is to be measured
  - Time-varying strain (ΔL/L): ~**10^{-21} /√Hz**
  - 5 pm/√Hz / 5 Gm
  - signal frequencies from 10^{-4} to 1 Hz,
  - signal durations of months to centuries

Constellation Response

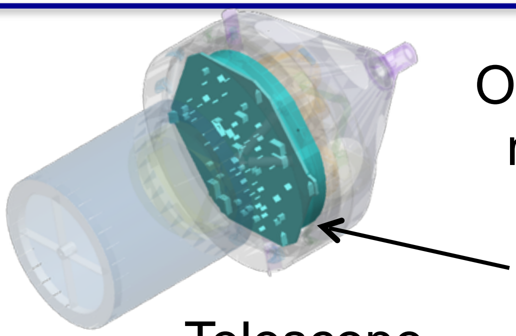


- Measurement concept
  - Measure distance changes between free-falling mirrors
  - Preferred measurement conditions:
    - A long measurement path to make ΔL large
    - A very quiet place to avoid disturbances to the test masses: SPACE!

Payload Integrated with Bus

Payload systems

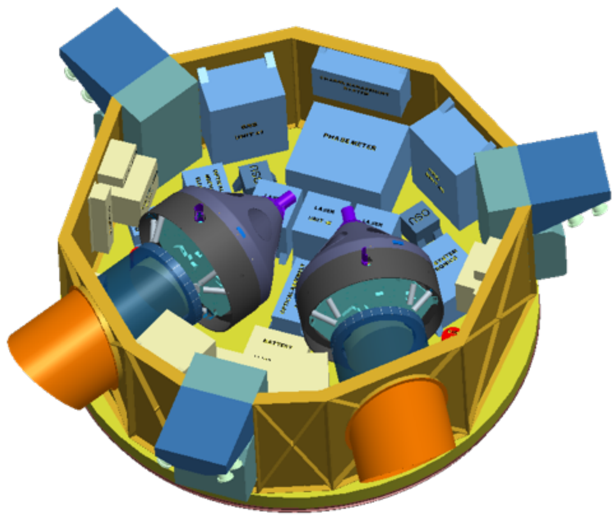
- Interferometer Measurement System (IMS)
  - Laser
  - Telescope
  - Optical bench
- Disturbance Reduction System (DRS)
  - Gravitational Reference Sensor (GRS)
  - μN thrusters
  - Control laws



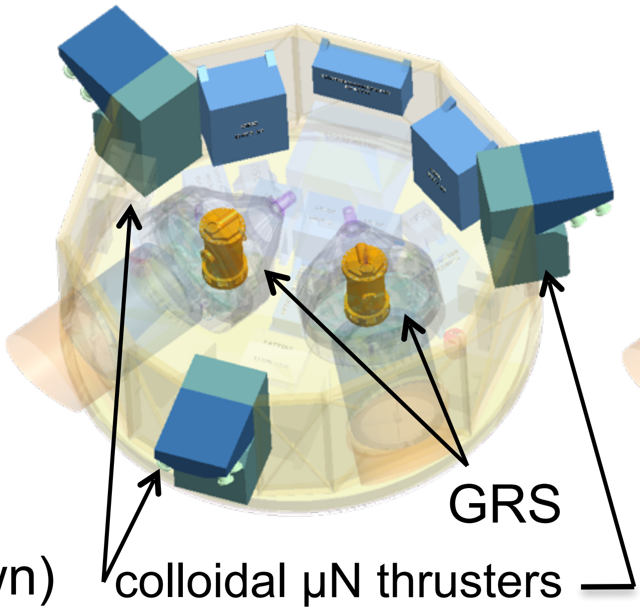
Optical bench mounted in Telescope Assembly

Telescope Assembly

Full Spacecraft Bus



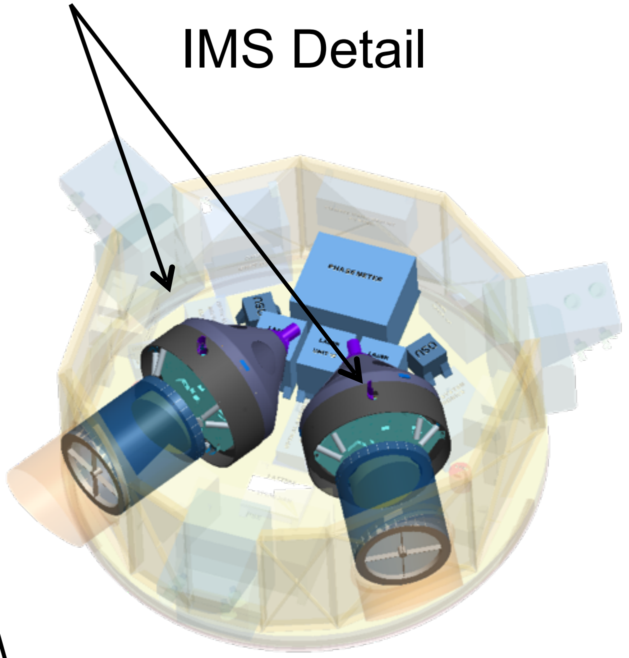
DRS Detail



GRS

colloidal μN thrusters

IMS Detail

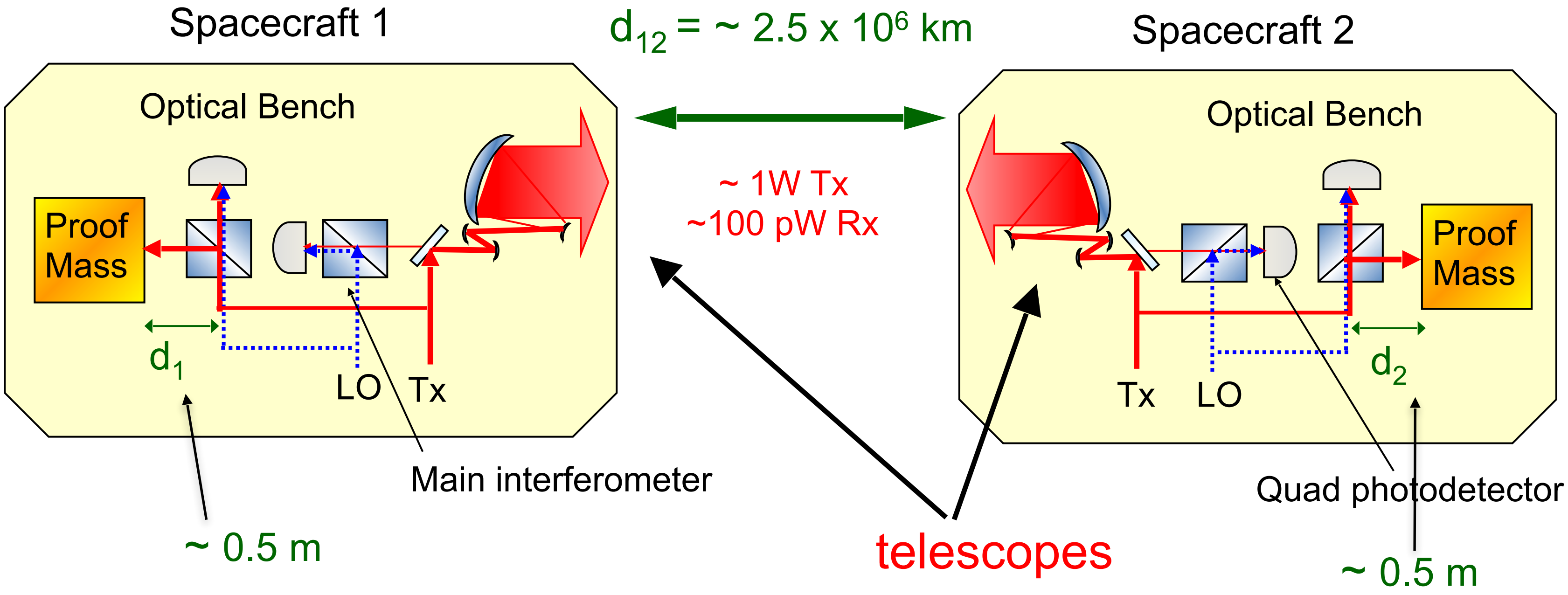


(Note: solar array not shown)

Inter-Spacecraft Distance Measurement

- **Test-mass to test-mass measured in 3 parts:**
  - 2 × test-mass to spacecraft measurements (short-arm: LPF tests this)
  - 1 × spacecraft to spacecraft interferometer (long-arm)

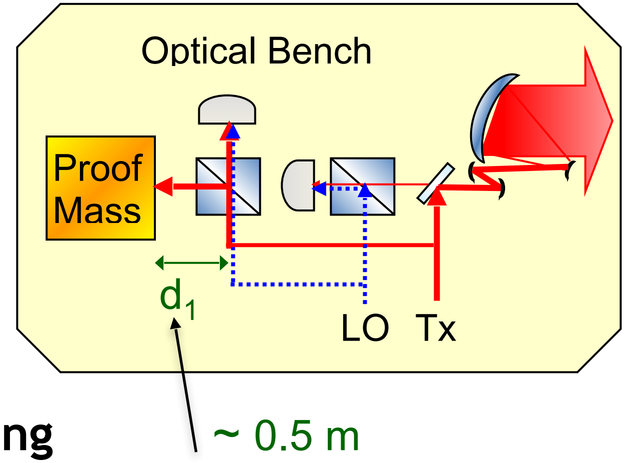
total separation = $d_1 + d_{12} + d_2$



telescopes

TELESCOPE DESCRIPTION

Telescope Functional Description/Requirements



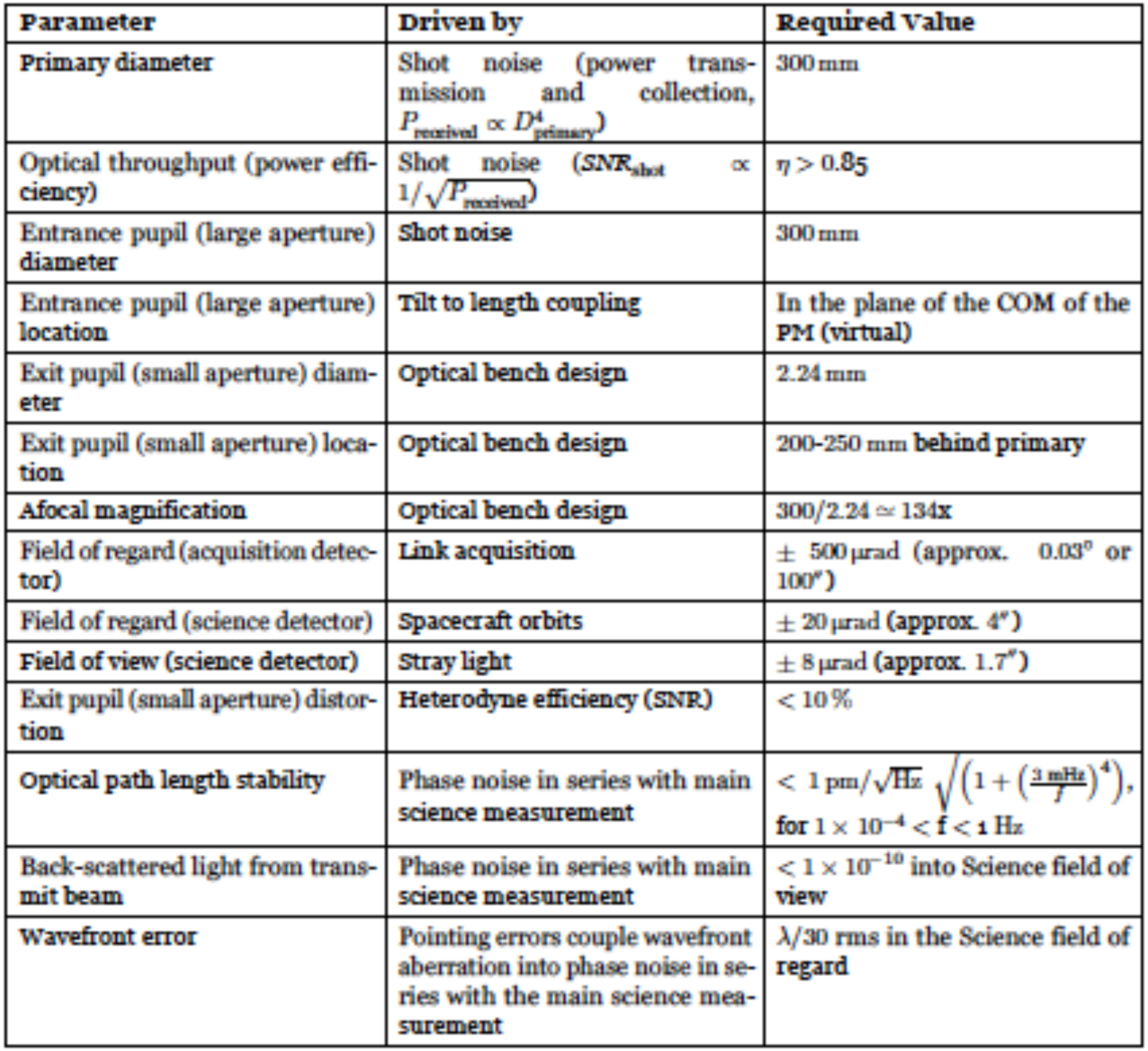
- **Afocal beam expander/reducer**
  - 300 mm dia. primary
  - 2.24 mm dia. on bench
  - 134X magnification
- **Simultaneous transmit and receive**

$$P_{\text{received}} \propto D^4_{\text{primary}}$$

- **Conjugate pupils to minimize tilt to length coupling**
  - Map angular motion of the spacecraft jitter to angular motion on the optical bench without lateral beam walk or piston
- **Smooth wavefront ($\lambda/30$) to minimize tilt to length coupling, also helps maximize on-axis power transmission**
- **Dimensionally stable (path-length fluctuations directly compete with pm scale measurement)**
- **Low back-scatter of transmit beam into receiver**

$$\sim 1\ \text{W transmitted}, \sim 500\ \text{pW received}$$

Key Telescope Requirements

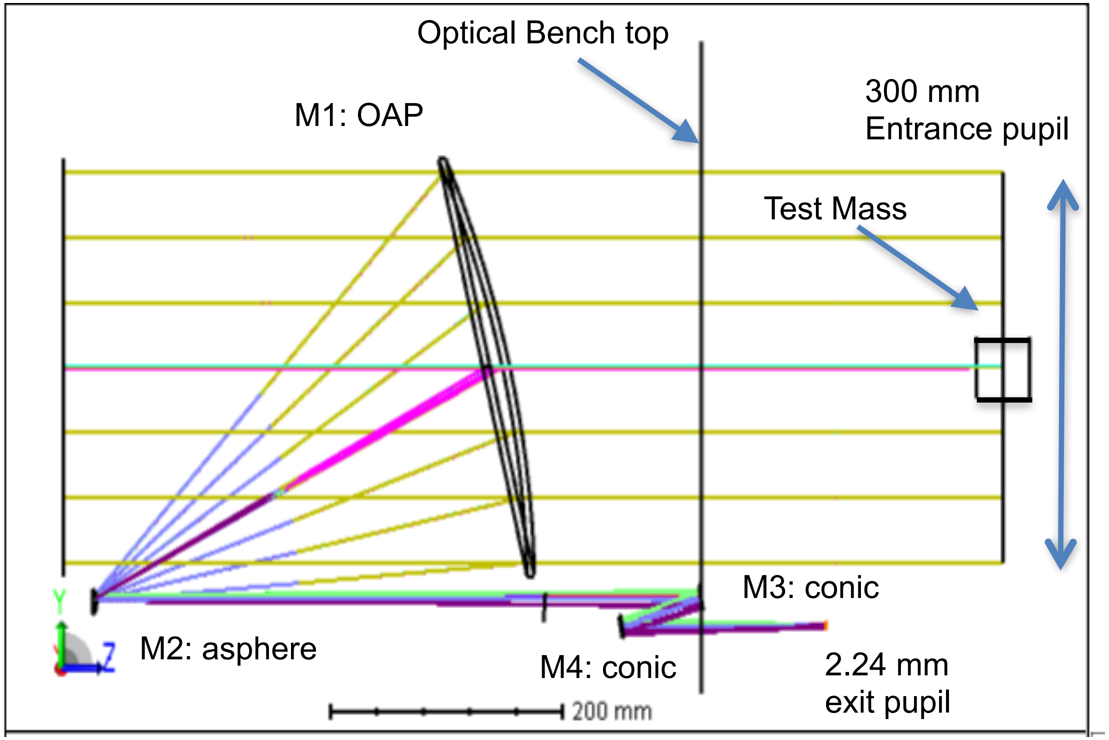


Parameter	Driven by	Required Value
Primary diameter	Shot noise (power transmission and collection, $P_{\text{received}} \propto D^4_{\text{primary}}$)	300 mm
Optical throughput (power efficiency)	Shot noise ($SNR_{\text{shot}} \propto 1/\sqrt{P_{\text{received}}}$)	$\eta > 0.85$
Entrance pupil (large aperture) diameter	Shot noise	300 mm
Entrance pupil (large aperture) location	Tilt to length coupling	In the plane of the COM of the PM (virtual)
Exit pupil (small aperture) diameter	Optical bench design	2.24 mm
Exit pupil (small aperture) location	Optical bench design	200-250 mm behind primary
Afocal magnification	Optical bench design	$300/2.24 \simeq 134$x
Field of regard (acquisition detector)	Link acquisition	± 500 μrad (approx. 0.03° or 100″)
Field of regard (science detector)	Spacecraft orbits	± 20 μrad (approx. 4″)
Field of view (science detector)	Stray light	± 8 μrad (approx. 1.7″)
Exit pupil (small aperture) distortion	Heterodyne efficiency (SNR)	< 10 %
Optical path length stability	Phase noise in series with main science measurement	$< 1\,\text{pm}/\sqrt{\text{Hz}}\,\sqrt{\left(1+\left(\frac{3\,\text{mHz}}{f}\right)^4\right)}$, for $1 \times 10^{-4} < f < 1$ Hz
Back-scattered light from transmit beam	Phase noise in series with main science measurement	$< 1 \times 10^{-10}$ into Science field of view
Wavefront error	Pointing errors couple wavefront aberration into phase noise in series with the main science measurement	$\lambda/30$ rms in the Science field of regard

challenging

challenging

Current 4-mirror Design



Optical Bench top

M1: OAP

300 mm
Entrance pupil

Test Mass

M3: conic

M2: asphere

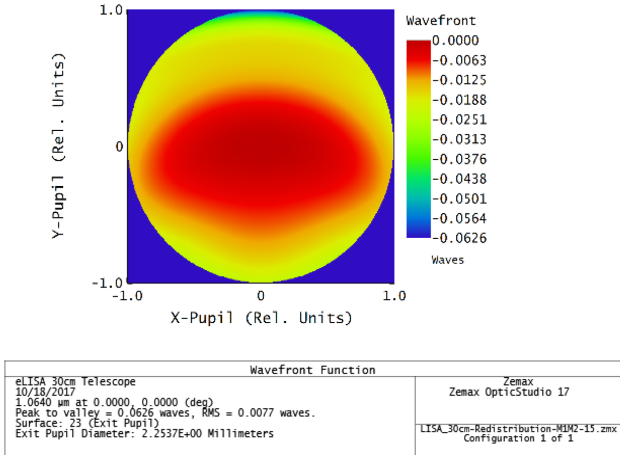
M4: conic

2.24 mm
exit pupil

M1/M2 Angular Magnification reduced from 74 to 55.8X (25% reduction)
M3/M4 now 2.4X, total is still 134X

Further M1/M2 Magnification reduction in process

Design residual WFE:
8.2 nm rms



Wavefront Function

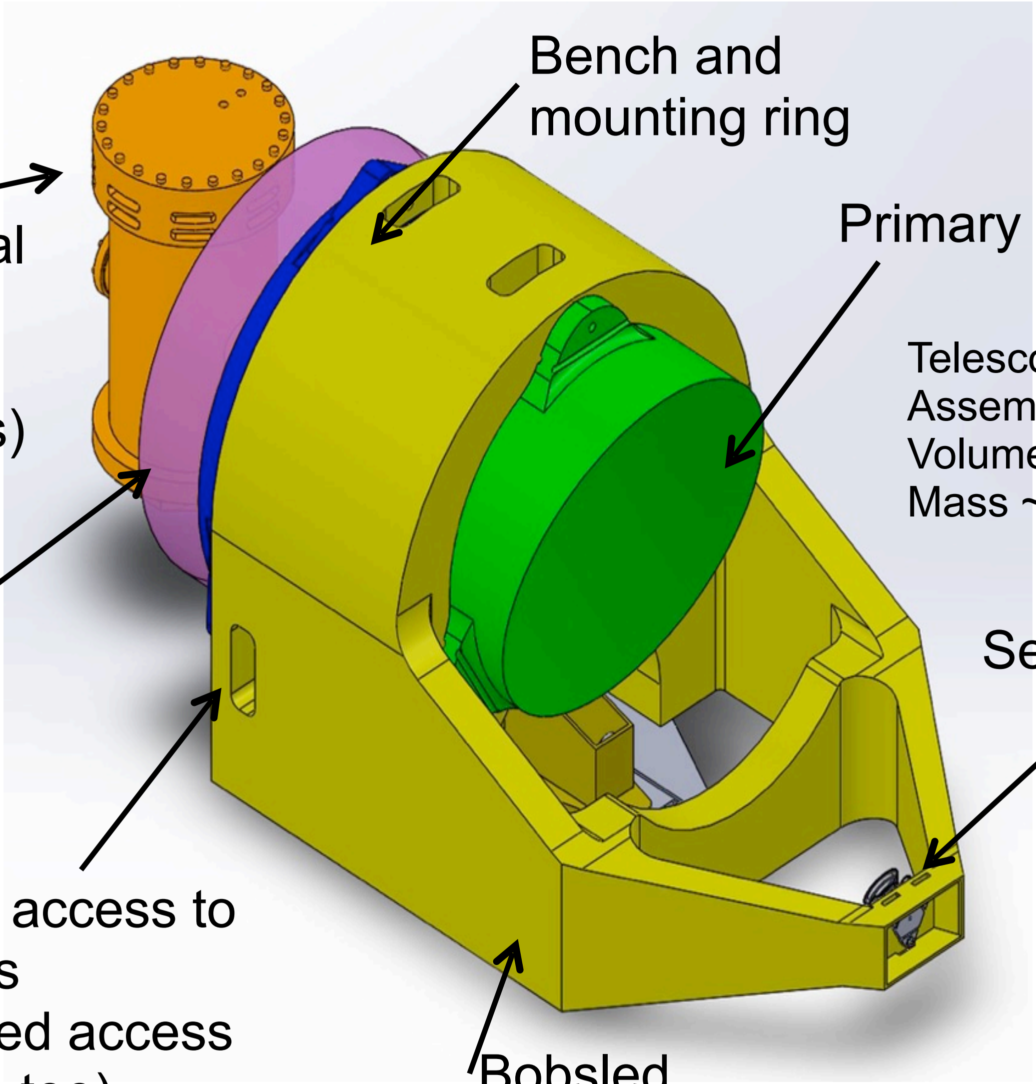
eLISA 30cm Telescope
10/18/2017
1.0640 µm at 0.0000, 0.0000 (deg)
Peak to valley = 0.0626 waves, RMS = 0.0077 waves.
Surface: 23 (Exit Pupil)
Exit Pupil Diameter: 2.2537E+00 Millimeters

Zemax
Zemax OpticStudio 17

LISA_30cm-Redistribution-M1M2-15.zmx
Configuration 1 of 1

- **Off-axis Cassegrain for stray light performance**
- **Schwarzschild-style pupil extender**
- **Simplified Design to reduce mirror cost, risk**

Extended "Bobsled"



Bench and mounting ring

Gravitational Reference Sensor (proof mass)

Primary

Telescope length ~ 450 mm
Assembly dia ~ 450 mm
Volume ~ 30 liters
Mass ~ 15 kg (just telescope)

Secondary

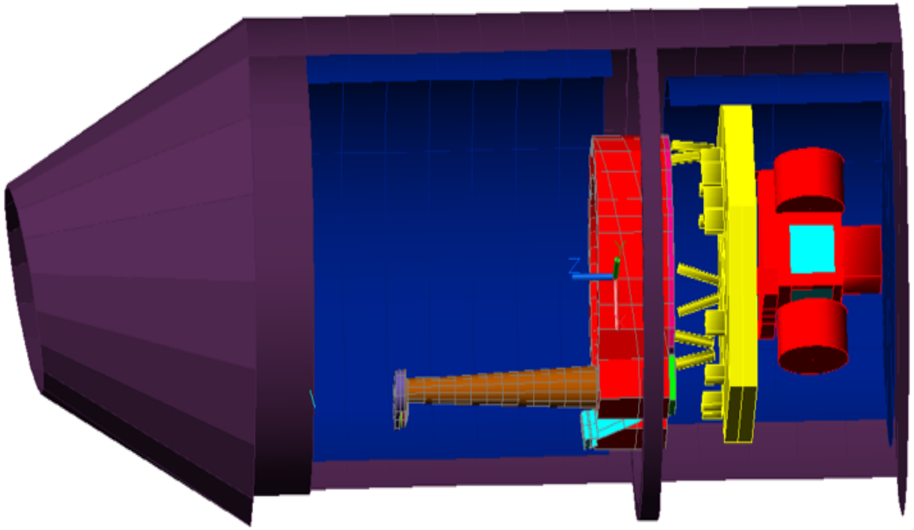
Rear "keep out" zone

Slots for access to fasteners (may need access to bench too)

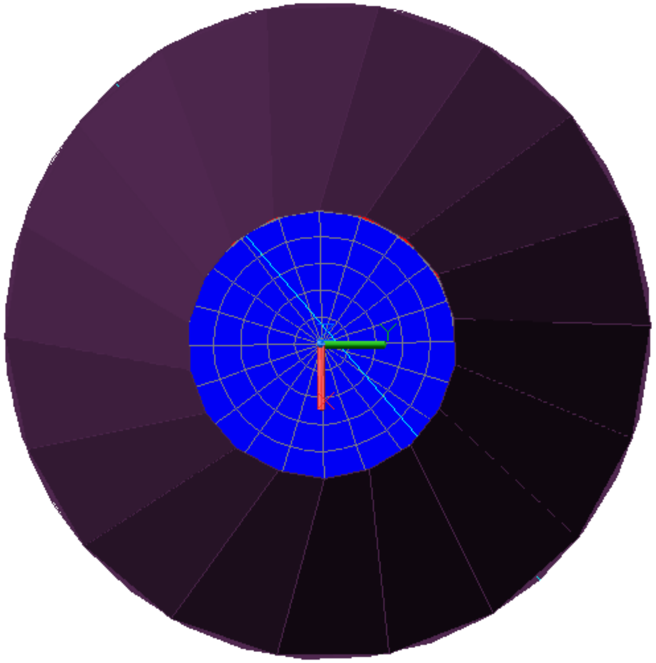
Bobsled

Note: this is a concept. Details are not finalized.

Preliminary Thermal Modeling



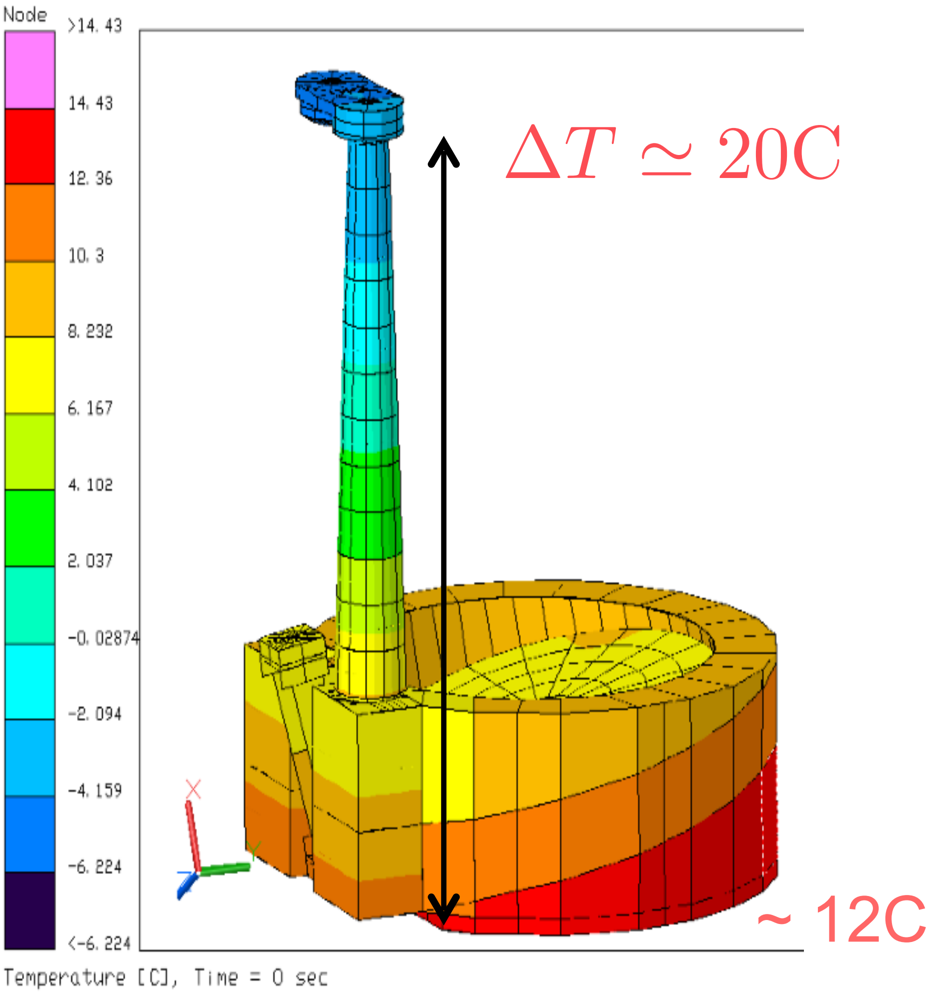
Primary baffled, secondary does not view cold space



View from space

Materials choice

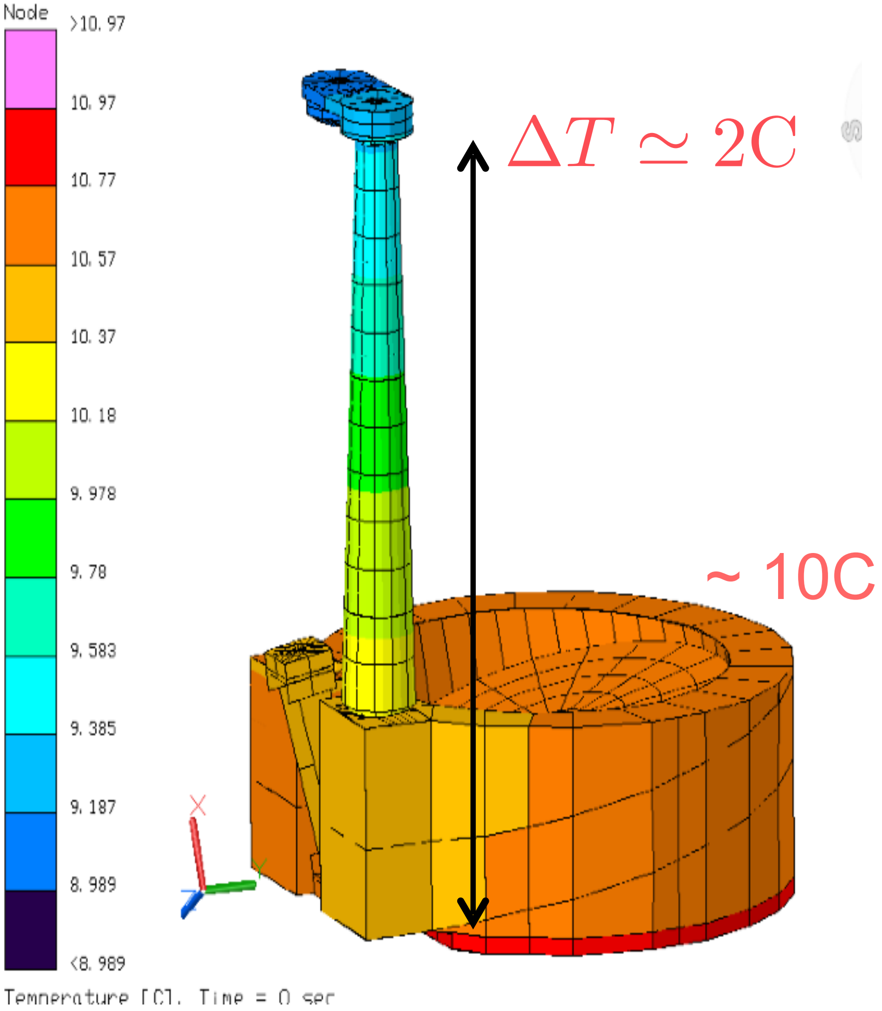
ZERODUR® like properties



$\Delta T \simeq 20\text{C}$

~ 12C

Silicon Carbide like properties



$\Delta T \simeq 2\text{C}$

~ 10C

CHALLENGES

SiC Spacer Dimensional Stability Demonstration

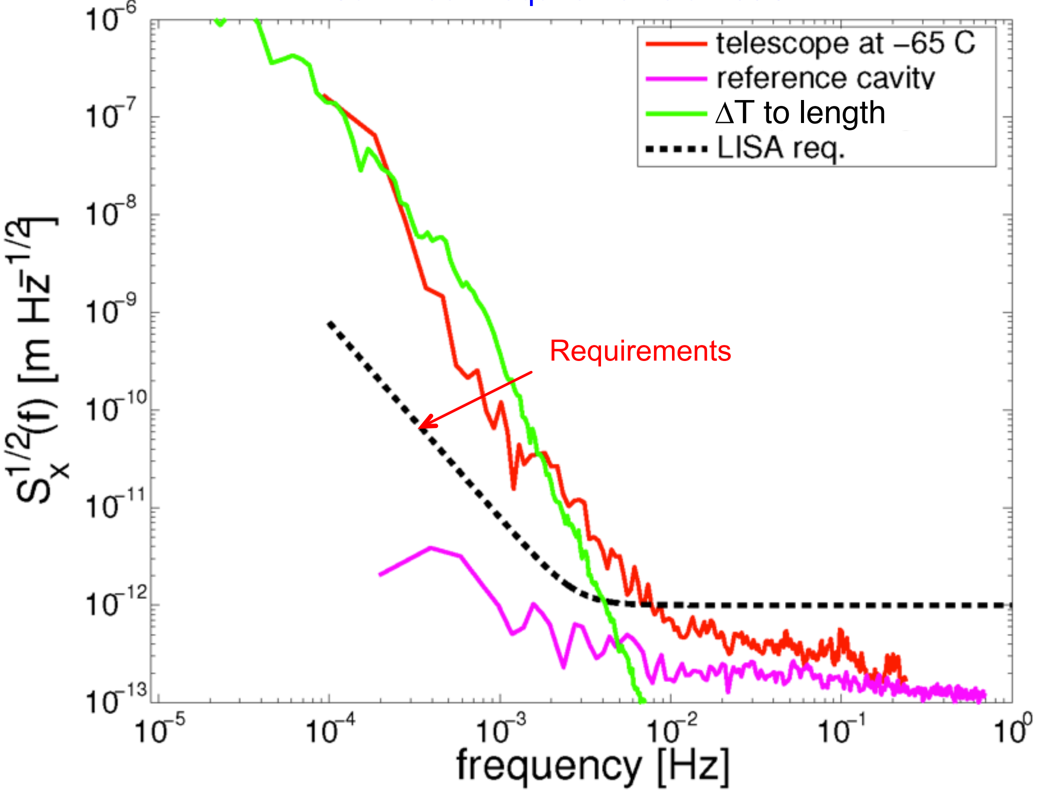
Spacer Activity Objective

- Develop and test a design for the main spacer element between the primary and secondary mirrors
- M1 - M2 spacing identified as critical by tolerance analysis
- SiC meets stability requirement with on-orbit ΔT(f)
- On-axis Quadpod would not meet scattered light requirement

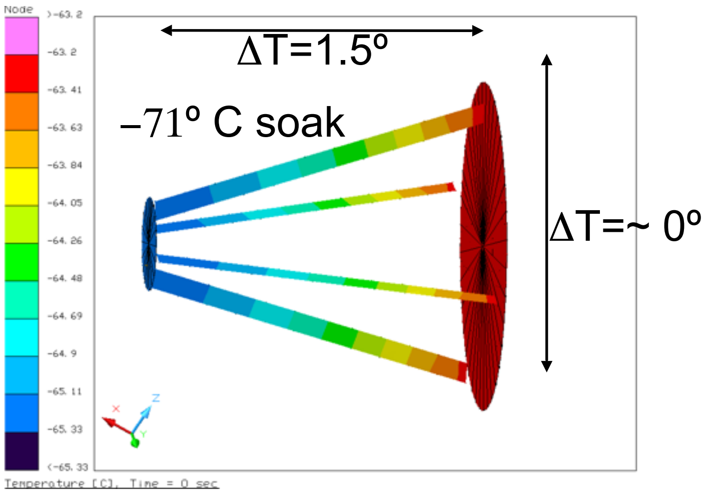
SiC Spacer Design: QuadPod

SiC Spacer Design

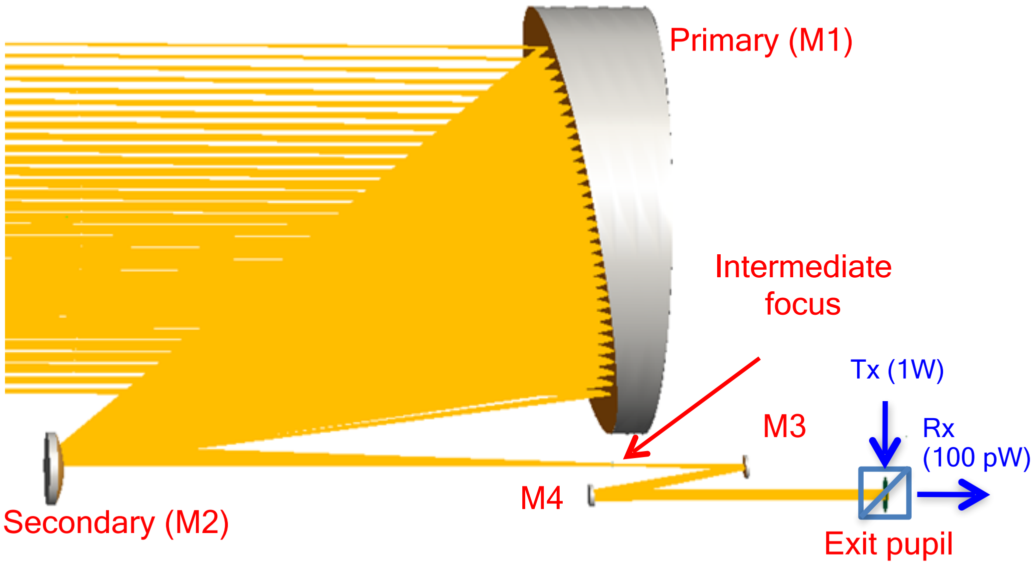
Can Meet Requirements at -65C

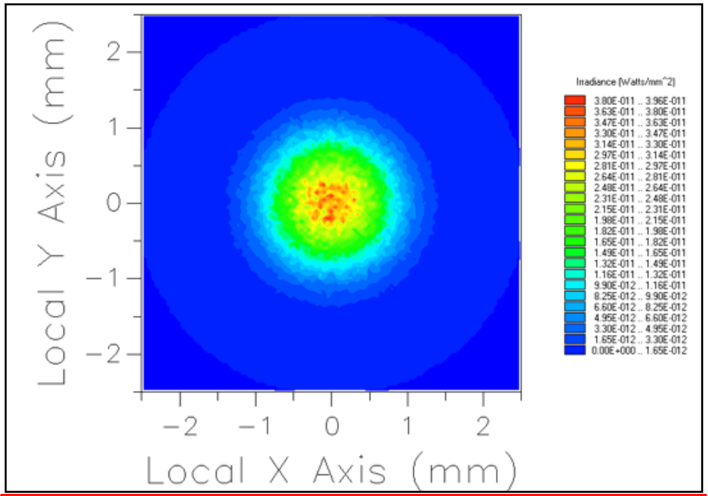


Thermal Model to Determine Test Conditions



Scattered Light Analysis

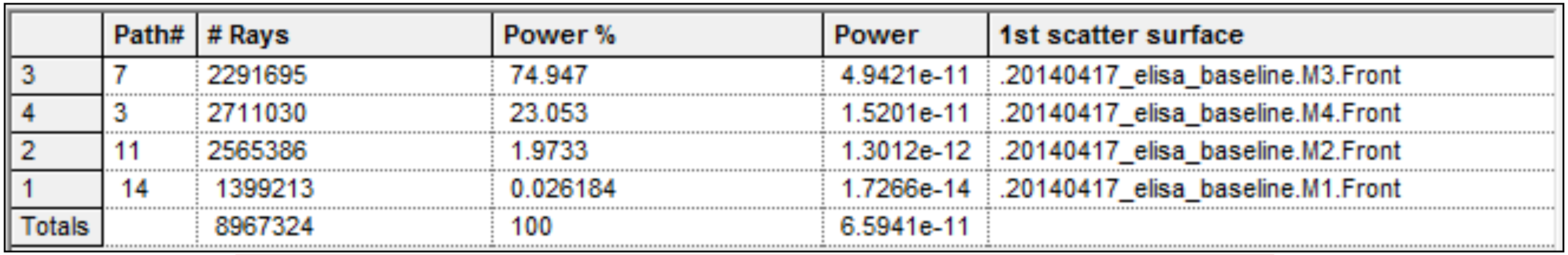



Pupil Plane Scatter Irradiance

Mirror	RMS surface roughness (Å)	MIL-STD 1246D CL
M1	15	300
M2	15	200
M3	5	200
M4	5	200

Conflicting accounts of on-orbit levels

- Source power = 1W
- Total power on the detector = 6.6×10^{-11} W → (barely) meets specification of less than 10^{-10}



	Path#	# Rays	Power %	Power	1st scatter surface
3	7	2291695	74.947	4.9421e-11	.20140417_elisa_baseline.M3.Front
4	3	2711030	23.053	1.5201e-11	.20140417_elisa_baseline.M4.Front
2	11	2565386	1.9733	1.3012e-12	.20140417_elisa_baseline.M2.Front
1	14	1399213	0.026184	1.7266e-14	.20140417_elisa_baseline.M1.Front
Totals		8967324	100	6.5941e-11	

aft optics contributes most of the scattered light

Summary

- **Gravitational waves enable dramatic new window on the Universe**
- **Precision metrology application drives requirements, not image quality**
  - Pico-meter-level pathlength stability
  - Low coherent backscattered light
  - Minimize tilt-to-length coupling
- **Requirements drive design**
  - Zerodur for pathlength stability
  - Off-axis for scattered light
  - Pupil relay to minimize tilt-to-length
- **Robust, manufacturable design**
  - Approximately 10 units needed